PHYSICAL PROPERTIES AND DETECTION OF SATURN’S LIGHTNING RADIO BURSTS

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Abstract

Physical properties of Saturn’s lightning radio bursts (occurrence, duration, flux, spectrum) are derived from Cassini-RPWS observations and compared to Voyager results. During the active storm of September 2004, peak occurrence rates, durations and spectral power were similar to those measured by Voyager. However, long-term occurrence appears very variable. Evidence is found for ionospheric propagation effects (cutoff and refraction), confirming the atmospheric origin of the observed bursts. The most intense ones should be detectable from the ground with presently available low-frequency radiotelescopes. Observation techniques are discussed together with preliminary results.

1 Lightning detection and occurrence

Nearly 24 years after the first Voyager-Saturn flyby which led to the discovery of Saturn’s lightning [Warwick et al., 1981; Kaiser et al., 1983], Cassini is performing long-term monitoring of these atmospheric emissions. We address here their physical properties as derived by the high frequency receiver (HFR) of the RPWS investigation [Gurnett et al., 2004]. Lightning radio emission consists of brief broadband bursts. When observed with a swept-frequency receiver this emission is detected as burst segments of apparent bandwidth Δf (the band swept during the occurrence of the burst) randomly distributed over the total true spectrum of the bursts. For the present study, automated detection over long intervals was performed by comparing each measurement with its immediate neighbours recorded with the same instrument status (frequency, antenna selection), and retaining measurements with intensity > x dB above these neighbours. In a second pass, fixed frequency interference and spurious broadband interference were automatically

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eliminated on the basis of their non random distribution in the time-frequency (t-f) plane. No ‘visual’ confirmation event by event was done here. The interval from 2004/01/01 (6 month prior to insertion in orbit) to 2005/03/31 is studied here, with a detection threshold $x = 1$ dB. We restricted frequencies to above 1.8 MHz (high-frequency band of the HFR). The RPWS/HFR is run in a number of specific modes of observation adapted to specific radio components. This powerful capability is also a difficulty when performing long-term statistics on the data. Due to the variable instrument setup, a relative threshold of 1 dB corresponds to a variable number of standard deviations of the background noise (thus to a variable effective S/N). Figure 1a illustrates this variable effective S/N threshold between $<2$ and $>5$ as a function of time over the period studied (the first 90 days of 2005 correspond to DOY 2004 367 to 457). Figure 1b shows the fraction of the time spent by the HFR dwelling at frequencies $>1.8$ MHz.

Figure 1: (a) Effective S/N corresponding to a burst detection threshold of 1 dB, versus time, in the high frequency band of the RPWS/HFR. Variations result from the various integration time and channel bandwidth used at different times. (b) HFR duty-cycle spent above 1.8 MHz. Dashed vertical lines delimitate the period of September 2004 studied in more details in this paper. The dotted vertical line marks 2005/01/01.

of events detected during the very active storm that took place from September 5th to 27th, 2004 (DOY 249 to 271). The detection threshold was quasi-steady at S/N $\sim 3.6$. The top panel shows the integrated occurrence per 15 min. intervals, which reveals a prominent periodicity of recurrence of the events (the computed period is 10h 40min, slightly shorter than the Saturn Kilometric Radiation period of 10h 45.75min measured
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by Cassini [Gurnett et al., 2005]). The bottom panel shows individual events in the t-f plane. These observations are consistent with a single storm system located at a latitude about 35° S, where the so-called ‘dragon storm’ was observed at the same time by the Imaging Science Subsystem on Cassini [Porco et al., 2005; Desch et al.]. The radio source rotates with Saturn’s atmosphere and bursts are observable once per turn during ‘episodes’ lasting for about half a rotation. Porco et al. [2005] noted a puzzling phase shift between the visible cloud and the radio burst episodes, the latter occurring on the average ∼ 0.2 Saturn rotation earlier than the former. Either the radio bursts must be produced some 70° ahead of the visible cloud, or they must be able to propagate from up to 70° beyond the horizon to Cassini. We propose elements for a solution below. Lightning occurrence is clearly very variable at timescales of days, revealing storm activity much more variable than at the time of both Voyager spacecraft [Zarka, 1985b], with sometimes totally missing episodes (e.g. on DOY 258). This occurrence variation is not due to the spacecraft-source distance, which remained between 125 and 150 Rs over the whole interval (1 Rs = Saturn’s radius ∼ 60300 km). An excess of spurious bursts is detected around DOY 264, related to an instrument setup change. The actual occurrence rate can be derived from the apparent one of Figure 2 via the HFR duty-cycle spent above 1.8 MHz, as plotted on Figure 1b. For the period studied here, this duty-cycle is quasi-permanently about 37%. The ∼ 4100 events detected above 1 dB on Figure 2 (excluding

Figure 2: (a) Burst occurrence per 15 min. intervals above a detection threshold of 1 dB, in September 2004. (b) Corresponding distribution of individual bursts versus time and frequency. Bursts are grouped in episodes of ∼ 5 hour duration, recurring at 10h 40min intervals.
those on DOY 264) translate thus in $\sim 11000$ actually emitted. The peak detection rate at $\sim 80$ events per 15 min. corresponds to one lightning flash produced every 4 sec, comparable to the occurrence rate derived by Voyager 1 at Saturn [Kaiser et al., 1983; Zarka and Pedersen, 1983]. The average rate, $\sim 4$ times lower, is comparable to that of Voyager 2 epoch.

To study in more detail the occurrence of events during an episode, we stacked on Figure 3 all the episodes of Figure 2b versus phase within Saturn's rotation. Both the apparition and disappearance of bursts seems to be frequency-dependent: lower frequencies start earlier than higher ones, and also stop earlier. Cassini was placed about 06:00 LT (05:55 to 06:30) during the entire interval studied, so that it could see the storm system emerge from Saturn's nightside and disappear under the dayside ionosphere. At disappearance,
the path from the source to the spacecraft crosses the ionosphere at an angle \( \approx \) equal to 
the storm longitude minus the spacecraft longitude \((\Lambda_{\text{storm}} - \Lambda_{s/c})\), so that radio waves 
sustain a low frequency (LF) cutoff at \( f_{\text{pe-peak}} / \cos(\Lambda_{\text{storm}} - \Lambda_{s/c}) \), plotted as the grey 
dashed curve on the right of Figure 3a. Let us remind that this cutoff results from 
reflection of radio waves coming from the atmosphere (where the radio refraction index 
is \( \mu \approx 1 \)) and entering the ionospheric plasma slab with peak plasma frequency \( f_{\text{pe-peak}} \) 
(where the refraction index for a wave frequency \( f \) is: \( \mu = (1 - f_{\text{pe-peak}}^2 / f^2)^{1/2} \)). This can 
explain (at least qualitatively) the frequency-dependent radio burst occurrence at the end 
of episodes [see also Zarka, 1985a]. The frequency-dependent apparition of the bursts (on 
the left of 3a) is more puzzling. We suggest that it is also due to a propagation effect, 
but this time through Saturn’s nightside ionosphere. The proposed effect comes from 
the refraction sustained by radio waves with frequency slightly above the cutoff: when it 
enters the ionosphere, the wave is refracted away from the normal to the plasma layer, and 
then back towards it when it exits. For a plane plasma slab, the net ray deviation is zero, 
but not for a spherical layer. The effect is dominant for rays entering the ionosphere at 
large angle from its local normal. Thus, we studied with 2D ray-tracing the propagation 
of rays launched horizontally from the storm source through an ionospheric layer with a 
Gaussian profile around a peak plasma frequency \( f_{\text{pe-peak}} \approx 1 \text{ MHz} \). If the ionosphere has 
spherical symmetry (i.e. no local time variation of its density) then the ray is either slightly 
deviated behind the limb (by \( \leq 1^\circ \), above a critical frequency equal to a few times \( f_{\text{pe-peak}} \)) 
or trapped under the ionospheric layer (below that frequency). However, if we introduce a 
decrease of the ionospheric electron density (and thus of \( f_{\text{pe-peak}} \)) versus local time in the 
nightside hemisphere, as can be expected from gradual recombination (see Figs. 6 and 7 
of [Moore et al., 2004] for Saturn’s case, and e.g. [Maruyama, 2002] for the Earth’s case), 
then rays can escape the ionospheric layer after temporary trapping, and thus experience 
an overall deviation of several degrees to several tens of degrees, as sketched in Figure 3b. 
This effect is stronger at lower frequencies (closer to \( f_{\text{pe-peak}} \)). Figure 3c: shows a series of 
results obtained by varying the ionospheric model parameters (average altitude and width 
of the Gaussian ionospheric layer, \( f_{\text{pe-peak}} \), and density decrease rate versus LT), where the 
ray deviation is translated in emersion phase of the radio burst source versus frequency, 
over the planet’s limb (located at phase 0.25). Although not accurately matching the 
observed emersion profile of Figure 3b, these curves obtained for a reasonable range of 
ionospheric parameters predict an effect consistent with the one observed, qualitatively 
as well as in order of magnitude. The above two propagation effects taken together 
can help solving the discrepancy between the radio burst occurrence and the visible cloud 
longitude mentioned earlier because they result in an overall phase shift of the radio burst 
ocurrence of the order of 0.1 rotation. Future detailed modelling may further improve 
the fit. Supporting the importance played by ionospheric propagation effects, we observed 
that the measured flux of lightning radio bursts is strongly attenuated (by a factor \( \sim 5 \)) 
at both episode edges compared to their flux at the center of the episodes.

2 Duration

RPWS/HFR measurements of September 2004 were made with a time constant of 45 msec 
between consecutively sampled frequencies. Lightning flash duration is directly derived
from the number of consecutive channels in which radio emission is detected. Figure 4 shows the distribution of these durations for events above a detection threshold of 1.5 dB. Flashes last for 1 to 8 consecutive channels, i.e. up to 360 msec. Their distribution is well described by an exponential law $N = N_0 e^{-D/D_0}$ with e-folding time $D_0 = 48$ msec (51 msec for 1 dB threshold). This result is very close to that obtained with Voyager data [Zarka and Pedersen, 1983; Zarka, 1985b]. Note that the extended radio spectrum of Saturn’s lightning (see below) implies that each flash probably consists of multiple short strokes, like terrestrial flashes [Farrell, 2000].

3 Lightning flux density and spectral power

Applying the absolute flux density calibration described by Zarka et al. [2004a], which relies on measurement by RPWS/HFR of the galactic background and careful estimation of the receiver background, we obtain the calibrated flux density $S$ of each detected burst (with an estimated accuracy better than $\sim 50\%$). Under the assumption of an isotropic emission, we derive the spectral power of the bursts $P = S \times 4\pi R^2$ with $R$ the spacecraft-source distance, taken here equal to the spacecraft-planet distance. Figure 5a shows the distribution of burst spectral power for the episodes of September 2004. The average spectrum is overplotted as a boldface line, and the average $+1\sigma$ (computed from the spectral power histogram at each frequency) as a lightface line. As with Voyager measurements, the spectrum is found to be nearly flat in the range 2-16 MHz, with an average power of $\sim 100$ W/Hz, thus a total power $\geq 1.6 \times 10^9$ W per flash. The strongest bursts may exceed 5 times this value. These numbers are within a factor 2 of independent flux and power estimates by Fischer et al. [2006].
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![Graphs showing spectral power, flux density, and NDA sensitivity versus frequency.](image)

Figure 5: (a) Spectral power versus frequency of September 2004 bursts, average spectrum (boldface line) and average +1σ (lightface line). The sky background contribution has been subtracted. (b) Corresponding flux densities at Earth (1Jy = 10^{-26}Wm^{-2}Hz^{-1}). (c) Same as (b) with different scales, superimposed on the 1σ detection threshold of \( \tau = 19 \) msec bursts of bandwidth \( b = 7 \) kHz or 10 MHz with the Nançay Decameter Array (NDA) connected to the R³ receiver (see text). The UTR-2 decameter array in Kharkov would have a detection threshold \( \sim 1 \) order of magnitude below NDA for the same \( b \) and \( \tau \). Earth’s ionospheric cutoff is indicated by the dotted vertical line.
4 Ground-based detection

From the above spectral power, the flux density at Earth $S_E$ is deduced as $S_E = P/4\pi R_E^2$ with $R_E$ the Earth-to-Saturn distance, $\sim 9.5$ AU on average. The flux densities corresponding to Figure 5a are displayed in Figure 5b,c, ranging from $\sim 200$ to $>1000$ Jy ($1\text{Jy} = 10^{-26}\text{Wm}^{-2}\text{Hz}^{-1}$). The possibility to detect these events using a ground-based LF radiotelescope results from the comparison of their flux to the level of sky background fluctuations, which is at 1$\sigma$ level:

$$\sigma_{\text{sky}} = 2kT_{\text{sky}}/A(b\tau)^{1/2}$$  \hspace{1cm} (1)

with $A$ the telescope effective area, $b$ the observation bandwidth, and $\tau$ the integration time. The LF sky background temperature can be approximated by:

$$T_{\text{sky}}(\text{K}) \sim 1.15 \times 10^8/f^{2.5}$$  \hspace{1cm} (2)

with $10 \leq f(\text{MHz}) \leq 100$ [van Haarlem et al., 2001]. For the Nançay Decameter Array (NDA) in France, the total area (adding both polarizations together because lightning radio emission is unpolarized) is $A = 48\lambda^2 \leq 7000$ m$^2$.

The resulting sensitivity is plotted versus frequency in Figure 5c: for $b = 7$ kHz and $b \sim 10$ MHz. The former bandwidth is one typical channel bandwidth of the new Nançay R$^3$ (Robust Radio Receiver of the Nançay Observatory [Weber et al., 2005]), while the latter is its typical total bandwidth (for $2000 \times 7$ kHz channels). With 10 MHz bandwidth, Saturn’s lightning should be detectable from Nançay above the Earth’s ionospheric cutoff $\sim 10$ MHz.

R$^3$ is a digital multichannel spectrograph with access to high (selectable) spectral and temporal resolutions and a large dynamic range. We have performed 5-hour observation runs of Saturn nearly every night from 2004/12/16 to 2005/02/03, with permanent tracking of Saturn, $\tau = 19$ msec time integration and a total band of 14 MHz divided in 2048 channels of $b = 7$ kHz bandwidth. The band 14-28 MHz was observed. Data was organized as t-f images of 5000 spectra ($\sim 1.5$ min.) $\times$2048 channels, and processed with a RFI-mitigation software developed in Meudon (RFI = Radio Frequency Interference). The algorithm, which is an improved version of the one described in Zarka et al. [1997], relies upon the fact that with the values of $\tau$ and $b$ chosen for the observations, Saturn’s lightning radio emissions are much weaker than sky background fluctuations (i.e. $S_E \ll \sigma_{\text{sky}}$).

Thus, any pixel with an intensity $\geq 3\sigma_{\text{sky}}$ above the background level is recognized as spurious. An iterative process is applied to compute $\sigma_{\text{sky}}$ in spite of the presence of RFI. Highpass filtering is applied during this process to eliminate slow signal fluctuations due to propagation through the Earth’s ionosphere. The spurious pixels recognized by the above method are then masked out, and the remaining ‘clean’ dynamic spectrum is integrated in frequency over the whole 14 MHz band (equivalent to a $\sim 10$ MHz RFI-free band), and also over the two intervals 14-21 MHz and 21-28 MHz separately (each equivalent to a $\sim 5$ MHz RFI-free band). This results in three series of integrated intensities versus time ($I_{14-28}(t)$, $I_{14-21}(t)$, and $I_{21-28}(t)$) corresponding to temporal integration $\tau = 19$ msec and effective bandwidth of 5 to 10 MHz. Corrections are applied for the spectral integration over a time variable number of clean channels (within each 19 msec spectrum), and
the final time series are normalized by their time variable standard deviation (over each 5000 × 2048 image). We performed tests suggesting that the algorithm is quite efficient, under the key condition that useful signal is ‘hidden’ in sky background fluctuations prior to spectral integration, which is indeed the case for Saturn lightning radio emission as shown in Figure d,b,c.

Unfortunately, Cassini observations revealed that lightning activity on Saturn was very low over the interval covered by these first Nançay observations. Close examination of RPWS dynamic spectra and occurrence diagrams suggests possible weak episodes between 2004/12/31 (DOY 2004 366) and 2005/01/03 (DOY 2004 369). Two nights of observations were performed at Nançay on DOY 366-7 and 368-9. Figure 6 shows the number of peaks > 3σ counted per 1.5 min. intervals (5000 values) in the time series integrated over the 14-21 MHz band. ~7 peaks are expected above 3σ for 5000 samples of pure random (Gaussian) noise. The higher number detected here may be due for a large part to weak residual RFI, and possibly to bursts of Saturnian origin. Direct comparison of Nançay and Cassini data is difficult because of the different local time and distance of the two observers relative to Saturn, Cassini being at 07:25 LT and Nançay at about 12:00 LT (and 67.4 light-minutes farther) on DOY 2004 366-369. Overall occurrence is expected to be shifted by ~ 5h42m between Cassini and Nançay, while intense bursts generated when the source is between these two LT might be seen by both observers. Examination of Cassini data suggests that part of the Nançay observation of DOY 366-7 and possibly the start of that of DOY 368-9 may have taken place at the time of Saturnian burst occurrence. On Figure 6, Nançay peak counts are plotted versus phase of Saturn’s rotation, in 0.1 rotation bins (~36° or 1 hour per bin). We remark that:

- the two observations sampled distinct phase ranges

![Figure 6: Number of peaks > 3σ per 1.5 min. intervals in the 14-21 MHz band detected with the Nançay R³ receiver during the nights of DOY 366-7 (open diamonds) and DOY 368-9 (filled diamonds), plotted versus phase of Saturn’s rotation (arbitrary origin). ~7 peaks > 3σ per 1.5 min. are expected for pure Gaussian noise. The solid line is the average detection rate over 0.1 rotation bins (~36° or 1 hour). Bins #1 and #6 contain very few observations and are thus subject to a larger error.](image-url)
• there appears to be an interval of 1-2 bins of low detection rate at the start of the first observation, followed by 3 bins of higher detection rate
• on the second night of observation, 1 ‘high’ bin is followed by 3 ‘low’ bins
• the whole data set organises in $\sim 5$ consecutive bins of high detection rate and 5 consecutive bins of low detection rate, consistent with observation of a single storm system rotating with the planet.

Although not yet fully conclusive, this result is encouraging.

5 Discussion and Conclusions

This study is still preliminary and a lot more will be done about Saturn’s lightning with RPWS until the end of the Cassini tour. However, the HFR has already proved to be a powerful lightning detector. Lightning radio bursts detected by Cassini have physical properties (short timescale occurrence, duration, flux, spectral power and spectrum shape) similar to those measured by Voyager 1 and 2 a quarter of a century ago. By contrast, the occurrence of lightning is found to be very variable at timescales of days to weeks [see also Desch et al.]. Voyager 2 observed an equatorial storm system similar to that discovered by Voyager 1 nine months earlier, only less well-defined or more spread in longitude. It was interpreted as being possibly the same, long-lasting storm system [Zarka, 1985b]. Cassini observes at intermediate southern latitudes well-defined storm systems appearing suddenly, lasting for days to weeks (during which the lightning source rotates with the atmosphere), with peak occurrence rates as high as $0.25 \text{ sec}^{-1}$, and then disappearing for weeks to months. The observations of September 2004 reveal strong ionospheric propagation effects through both the nightside and dayside ionosphere. The detailed modelling of burst occurrence and intensity versus frequency and geometry of observation should eventually bring new information on Saturn’s ionosphere. But the mere existence of these effects confirms – if it were necessary – the atmospheric origin of the detected radio bursts.

Intense events appear to be detectable from the ground even by medium-size instruments as the NDA (taking advantage of the broad intrinsic bandwidth of the emission and of modern RFI-mitigation capabilities), and a fortiori by larger instruments as the UTR-2 array in Kharkov [Sidorchuk et al.; Zarka et al., 1997] and soon LOFAR [Zarka et al., 2004b]. We have obtained encouraging but not yet conclusive results: high detection rates in Nançay (higher than from the more sensitive RPWS) suggest that only a small fraction of the events detected from the ground may be of Saturnian origin, but these events show a coherent organization versus phase of Saturn’s rotation. This first campaign of ground-based observations took place during a period of low lightning activity on Saturn, and it will be repeated at the next occurrence of intense activity not too far from Saturn’s opposition (daytime observations suffer a too high level of RFI pollution).

We will also attempt to locate directly the instantaneous sources of the most intense bursts (with $S/N > 15 - 20 \text{ dB}$) in the atmosphere using the so-called ‘direction-finding’ capability of RPWS/HFR [see Cecconi et al.]. Finally, a major question to be addressed
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by RPWS is the existence of Titan lightning. Preliminary analysis of the first 6 Titan flybys between 2004/01/01 and 2005/03/31 did not reveal any strong event. A few flybys were polluted by solar type III emissions, which should be eliminated before searching for statistical evidence of weak Titan lightning (as $1/R^2$ occurrence or intensity variation). The good sensitivity of RPWS exploited during the 40 Titan flybys to come should allow to answer this question. Besides, study of the radio emission of Saturn lightning is expected to provide significant advances on lightning physics, atmospheric and ionospheric physics, and be an ideal testbed for the validation of LF ground-based observations at high sensitivity.

References


